Dark matter

Experimental evidence for the Big Bang model is:

- A nearly uniform distribution of matter in the universe.
- The universe expands.
- The cosmic background radiation which has a temperature of 2.7 K (0.0002 eV).
- An abundance of light elements (He, D, Li)

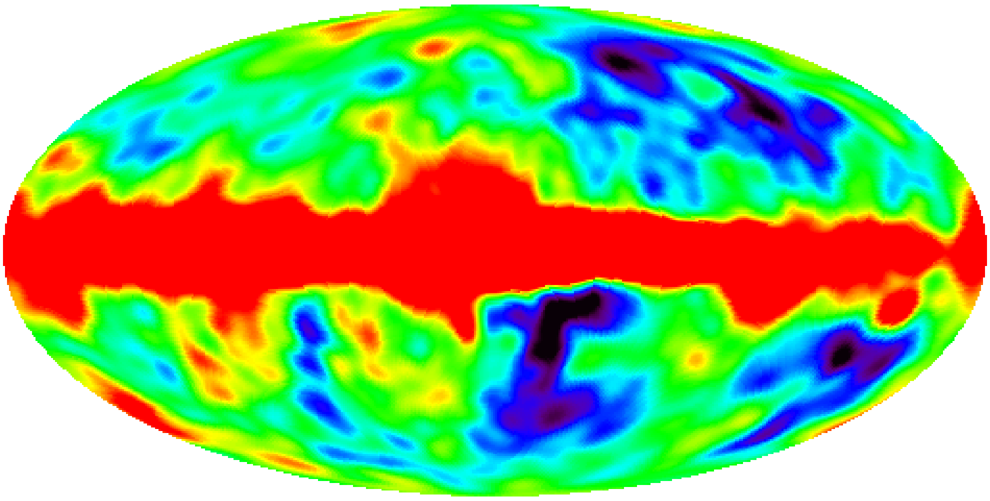

Figure 141: Sky as seen at microwave frequencies by the COBE satellite. Red (hottest) and blue (coldest) regions differ by only 0.0002 K while the overall temperature is 2.7 K

– If the density of the universe is smaller than the critical density, the expansion of the universe will continue for ever.

The critical density: $$\rho_c = \frac{3H_0^2}{8\pi G} = O(10^{-26})\ kg\ m^3$$

Where H_0 is the Hubble constant and G is the gravitational constant.

In the inflationary Big Bang model, the density of the universe is estimated to be close to the critical density:

$$\Omega \equiv \rho / \rho_c = 1$$

Where Ω is called the relative density. However, the observable (i.e. emitting electromagnetic radiation) matter in the Universe give only $\Omega_L \approx 0.01$

→ The rest is called "dark matter"

Possible components of the dark matter:

a) *Baryonic matter* that emit little or no e.m. radiation: brown dwarfs, small black holes – MACHO's (for MAssive Compact Halo Object). There is evidence that $\Omega_B \approx 0.06$ only.

b) If neutrinos have a mass > 1eV they would make a significant contribution to the density of the universe ("*hot dark matter*") . It is, however, difficult to explain how the galaxies have formed if neutrinos are the dark matter.

c) "*Cold dark matter*": WIMP's (Weakly Interacting Massive Particles), non-baryonic objects, non-relativistic at the early stages of the evolution of the universe.

The search for WIMPs

❖ Interactions between WIMPs and matter are very rare. About one WIMP per day is expected to interact in each kg of matter.

❖ To minimize the background, the WIMP detectors are installed deep underground and surrounded with shielding.

❖ The Boulby experiment uses a NaI detector which produces scintillation light if a WIMP interacts with an atom. 200 tons of ultra pure water is used as shielding.

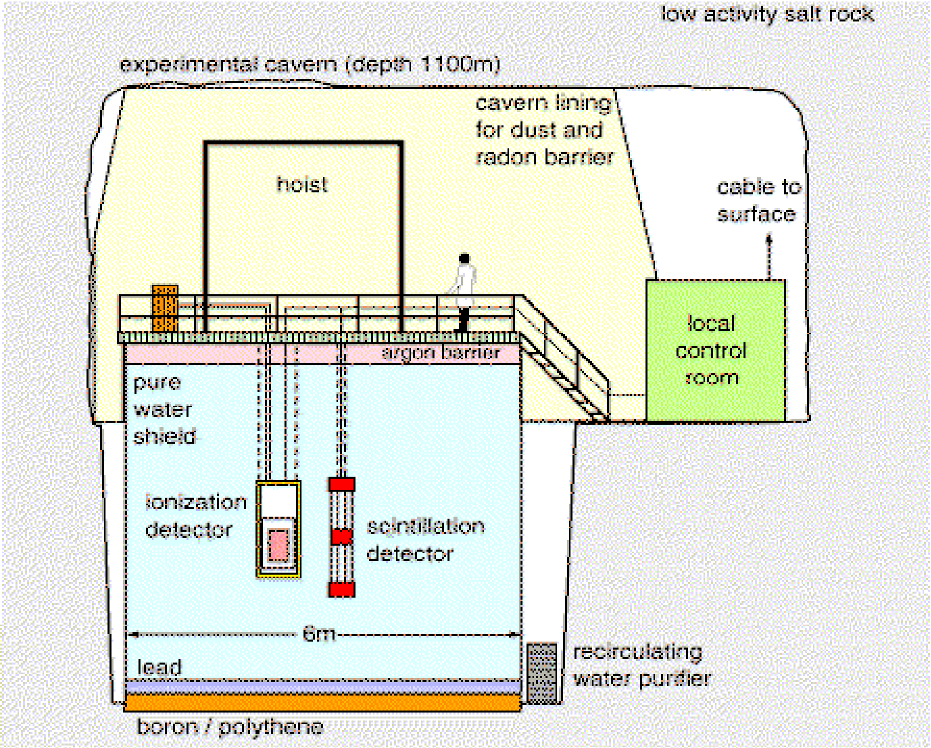

Figure 142: Layout of the Boulby experiment in the UK.

Grand Unified Theories (GUTs)

- Weak and electromagnetic interactions are unified, why not to add the strong one?
- At some very high “unification mass” electroweak and strong couplings might become equal

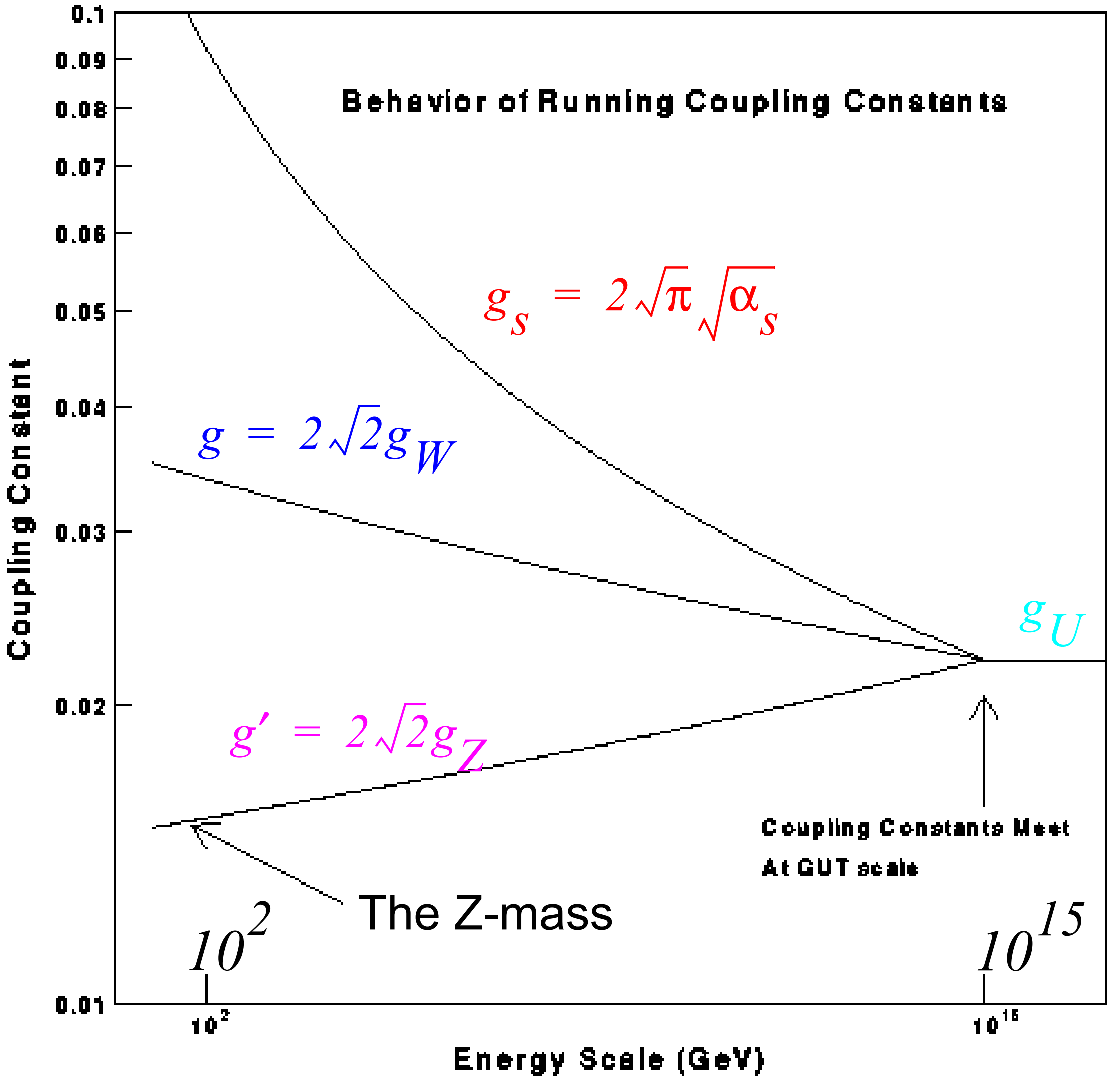

Figure 143: Behavior of the coupling constants in GUT

Grand unified theories can be constructed in many different ways.

◆ The Georgi-Glashow model combines coloured quarks and leptons in single families, like

$$(d_r, d_g, d_b, e^+, \bar{\nu}_e)$$

and hence new gauge bosons appear:
X with Q=-4/3 and Y with Q=-1/3, $M_X \approx 10^{15}$ GeV/c^2:

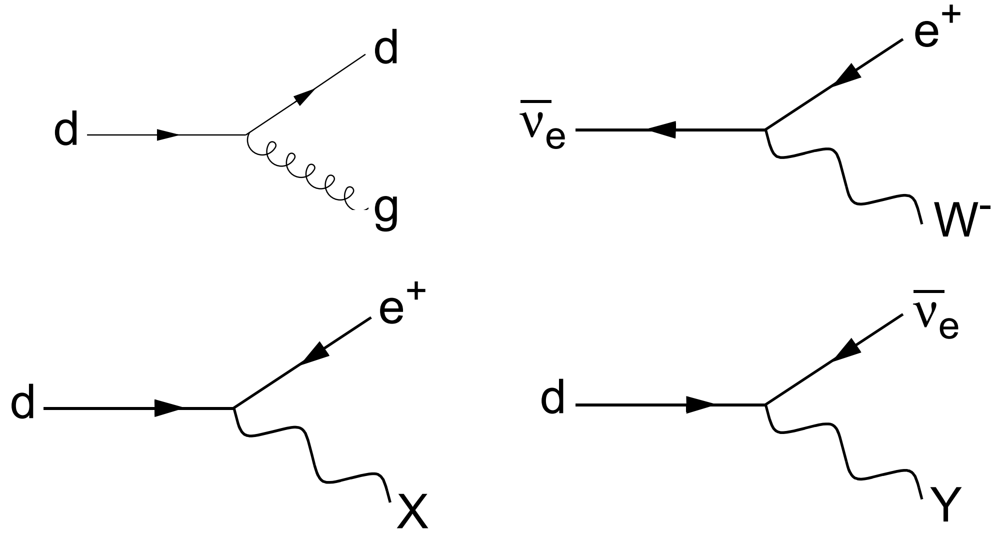

Figure 144: Standard processes together with new ones predicted by GUT

The single unified coupling constant is g_U, and

$$\alpha_U \equiv \frac{g_U^2}{4\pi} \approx \frac{1}{42} \qquad (166)$$

→ The Georgi-Glashow model explains why the electron and the proton have the same charge

According to the model the sum of electric charges in any given family must be zero ⇒ $3Q_d + e = 0$ ⇒ the down-quark has charge -e/3.

❖ The factor of 3 arises simply from the number of colours

→ This model also predicts the weak mixing angle since it predicts the value of one of the three coupling constants:

$$\sin^2\theta_W = 0{,}21 \tag{167}$$

This is close to the measured value of the weak mixing angle.

Proton decay

GUT predicts that the proton is unstable and that it can decay by a process involving X or Y bosons

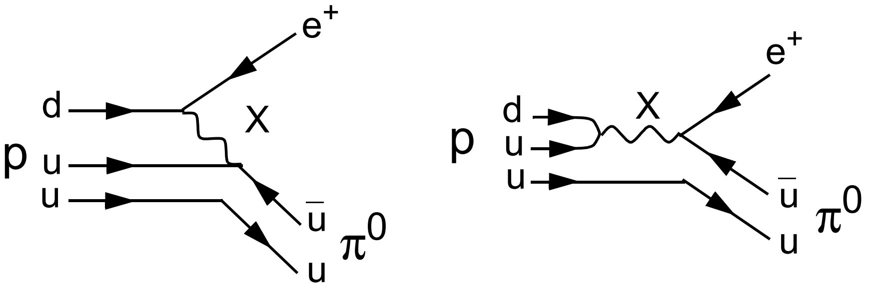

Figure 145: Proton decays in GUT

In processes like those above, baryon and lepton numbers are not conserved, but the combination

$$B - L \equiv B - \sum_{\alpha} L_{\alpha} \quad (\alpha = e, \mu, \tau) \tag{168}$$

is conserved.

❖ From a simple zero-range approximation, the lifetime of the proton can be estimated to be:

$$\tau_p = 10^{29} \div 10^{30} \; years \tag{169}$$

while the age of the universe is only about 10^{10} years...

- Some detectors which are used in neutrino physics (IMB, Kamiokande) are also looking for the proton decays.

- The most looked for decay mode is

$$p \rightarrow \pi^0 + e^+ \rightarrow \gamma\gamma + e^+$$

where the experiments looks for one positron and two electron-photon pairs from photon conversions.

- No clear examples of proton decays have been observed and the upper limit on the proton lifetime is now:

$$\frac{\tau_p}{B(p \rightarrow \pi^0 e)} > 5 \times 10^{32}\ years$$

- The Georgi-Glashow model predicts this ratio to be only 0.003×10^{32} - 0.03×10^{32} years in disagreement with the experiments. Other GUT models, however, predicts longer lifetimes.

The cosmic baryon asymmetry

❖ Why are there more baryons than antibaryons in the universe ?

Answer:

1. There was always an excess of baryons (the baryon number is conserved).

2. At the time of the Big Bang the universe had zero baryon number. The baryons were produced later (the baryon number is not conserved as suggested by GUT).

In the second case it is also necessary that C and CP are not conserved so that more antiparticles can be transformed to particles than vice versa.

Supersymmetry (SUSY)

→ The most popular GUTs incorporate supersymmetry (SUSY).

- Every known elementary particle has a supersymmetric partner -"superparticle" - with different spin:

Particle	Symbol	Spin	Superparticle	Symbol	Spin
Quark	q	1/2	Squark	$\tilde{q}$	0
Electron	e	1/2	Selectron	$\tilde{e}$	0
Muon	μ	1/2	Smuon	$\tilde{\mu}$	0
Tauon	τ	1/2	Stauon	$\tilde{\tau}$	0
W	W	1	Wino	$\tilde{W}$	1/2
Z	Z	1	Zino	$\tilde{Z}$	1/2
Photon	γ	1	Photino	$\tilde{\gamma}$	1/2
Gluon	g	1	Gluino	$\tilde{g}$	1/2
Higgs	H	0	Higgsino	$\tilde{H}$	1/2

Supersymmetric particles have to be much heavier than their counterparts since they are not observed.

- SUSY shifts the grand unification mass from 10^{15} to 10^{16} GeV/c^2, and hence the lifetime of the proton increases:

$$\tau_p = 10^{32} \div 10^{33} \text{ years} \qquad (170)$$

which is more consistent with experimental (non)observations.

- SUSY also predicts a value of the weak mixing angle which is closer to the experimental results.

- SUSY models even attempts to unify ALL forces, including gravity, at the *Planck mass* of order 10^{19} GeV/c^2 by replacing particles with *superstrings*

- The lightest superparticles can be candidates for the cold dark matter. Most models introduce a *neutralino* $\tilde{\chi}_0$, which is a mixture of photino, Higgsino and zino.

→ One possibility to look for SUSY at LEP is to search for selectron production followed by a decay to electrons and neutralinos:

$$e^+ + e^- \to \tilde{e}^+ + \tilde{e}^-$$

$$\tilde{e}^+ \to e^+ + \tilde{\chi}_0 \qquad \tilde{e}^- \to e^- + \tilde{\chi}_0$$

1) The cross section for producing selectron pairs is comparable with that of producing ordinary charged particles of the <u>same</u> mass

2) The selectrons decay before they can reach a detector

3) Neutralinos are virtually undetectable due to very weak interaction

The events one is looking for has only final state electrons and these

a) carry only about half of the collision energy

b) are not emitted in the opposite directions in the centre-of-mass frame

- No events with a neutralino signature have been observed.

- A measurement by DELPHI, using many more searches than slepton searches, set a lower limit on the neutralino mass of 37 Gev:

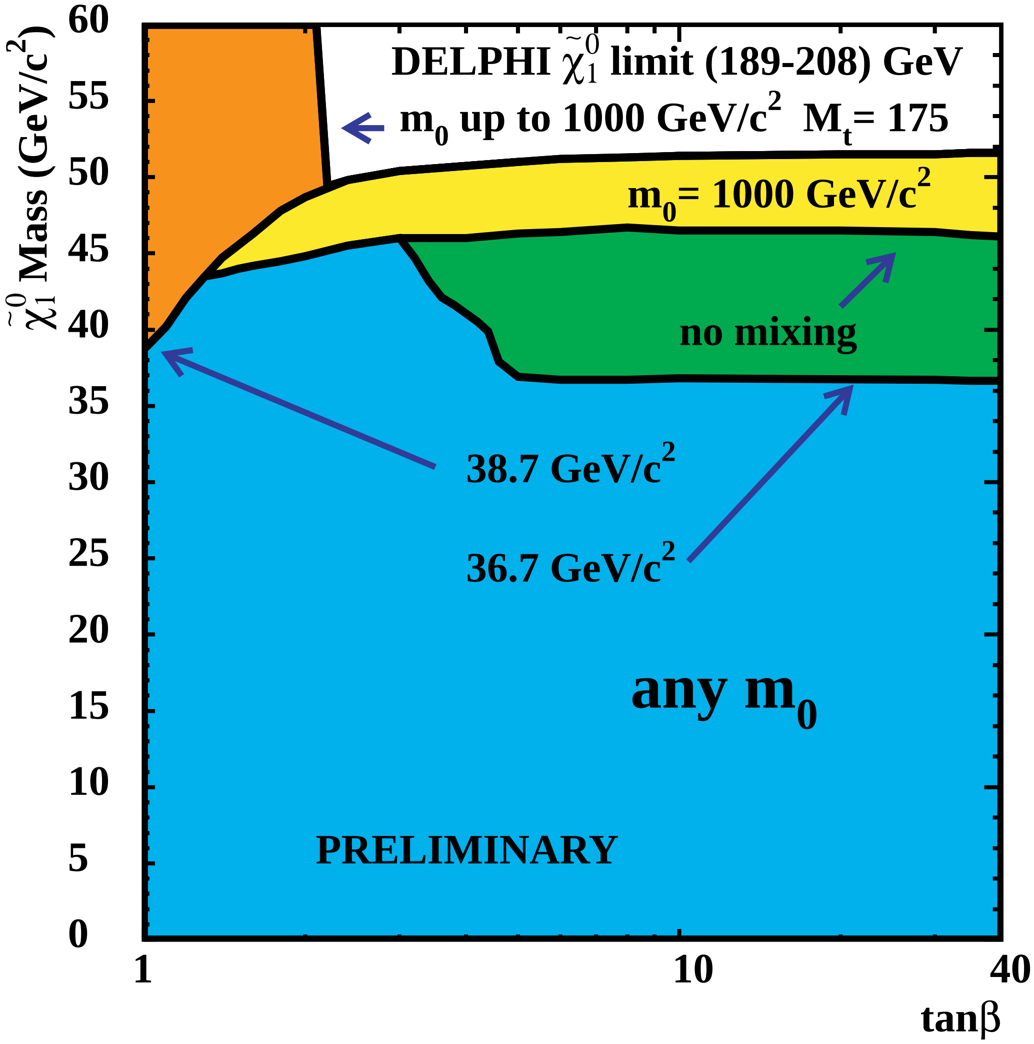

Figure 146: The lower limit on the mass of the lightest neutralino as a function of *tan* β (the ratio of the vacuum expectation values of the two SUSY Higgs doublets). m_0 is a universal SUSY mass parameter of the sfermions.

Summary

• Neutrinos

a) Neutrino mixing

b) Neutrino oscillations

c) Methods to detect neutrino oscillations

d) The atmospheric neutrino anomaly

e) The solar neutrino problem

• Dark matter

f) What is dark matter ?

g) Candidates for dark matter

• Grand Unified Theories

h) All coupling constants equal

i) The Georgi-Glashow model

j) The importance of proton decay

• Supersymmetry

k) Superparticles with different spin

l) Unification of all forces including gravity

m) The search for neutralinos